

Peculiar Hubble Flows in Our Local Universe

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ABSTRACT

A formalism that simultaneously searches for the monopolar and dipolar peculiar velocities is presented. The formalism is applied to (1) the Mark III catalogue, (2) Lauer and Postman’s Abell cluster catalogue, and (3) Riess et al.’s Type Ia supernova sample. The emphasis is drawn to the monopolar peculiar velocities, i.e., peculiar Hubble flows, within these samples. The samples show inconsistent peculiar Hubble flows within a depth of $\sim 60h^{-1}$ Mpc. Beyond a depth of $\sim 80h^{-1}$ Mpc, the Hubble flows of all samples converge to the global Hubble flow to better than 10% at the 2σ level. The results are compared with theoretical predictions. They at face value disfavor models predicting smaller peculiar velocities such as the tilted Cold Dark Matter model. Limitations of the catalogues are discussed, so are ways to improve the catalogues so that an accurate map of Hubble flows in our local universe can be drawn and be compared with theoretical predictions.

Subject headings: cosmology: observations — cosmology: theory — Large-scale structure of universe

1. Introduction

Peculiar Hubble flows (monopolar deviations from a global Hubble flow) are as important as bulk motions (dipolar deviations from a global Hubble flow) in reflecting the underlying density fluctuation of the universe. But they have not received the same attention as bulk motions have, because the uncertainty of our knowledge about the true Hubble constant H_0 is often larger than the expected peculiar Hubble flows at large scales. One can, however, investigate the variation of Hubble flows within a sample without knowing the value of H_0 . Moreover, the inferred Hubble flows from samples that extend significantly beyond $100h^{-1}$ Mpc should be so close to the global Hubble flow that it is meaningful to investigate the peculiar Hubble flows within the samples. In doing so, the ultimate goal is to map out the variation of the Hubble expansion as a function of depth, which is directly related to the underlying density fluctuation (Shi, Widrow and Dursi 1996; Turner, Cen and Ostriker 1992), and then to use the peculiar Hubble flows detected to test theoretical models.

Here I first present a formalism that simultaneously calculates the peculiar Hubble flows and the bulk motions, along with their errors, within a sample. The formalism is applied to (1) the Mark III catalogue based on the Tully-Fisher relation as a distance indicator (Willick et al. 1996), (2) The Abell cluster catalogue of Lauer and Postman (1994, LP thereafter) with the brightest cluster galaxies (BCGs) as distance indicators, and (3) the sample of 20 Type Ia supernovae (SNe) of Riess et al. (Riess, Press, Kirshner 1996). All three samples are deep or fairly deep: the Mark III catalogue extends to 11,000 km/sec; the LP sample is complete up to 15,000 km/sec; the SN sample probes the deepest, 28,000 km/sec. In other words, Hubble flows defined by these samples should approximate the global Hubble flow to a fairly high degree (see table 1 and text).

The results from the three samples are presented and analyzed in section 3–5. They

are compared with theoretical predictions from four representative models (table 1): the standard Cold Dark Matter (sCDM) model, a tilted Cold Dark Matter (tCDM) model, a vacuum energy Λ dominated Cold Dark Matter (Λ CDM) model, and a Cold plus Hot Dark Matter (CHDM) model. Finally I conclude in section 6 and comment briefly on improving the measurements of peculiar Hubble flows and their potential to test theoretical models.

2. Formalism

The peculiar Hubble flow and the bulk motion of a sample can be obtained by maximizing the likelihood (Kaiser 1988, 1991)

$$L(U_i, \delta H) = \prod_q \frac{1}{\sigma_q^2} \exp\left[-\frac{(S_q - \delta H r_q - U_i \hat{r}_q^i)^2}{2\sigma_q^2}\right], \quad (1)$$

where $\mathbf{r}_q (= r_q^i \hat{\mathbf{r}}_i, i = x, y, z)$ is the position of an object in the sample, S_q its estimated line-of-sight peculiar velocity with an uncertainty σ_q . $\mathbf{U} (= U_i \hat{\mathbf{r}}_i, i = x, y, z)$ is the bulk motion of the sample, and δH is its peculiar Hubble expansion rate. Maximizing $L(U_i, \delta H)$ with respect to U_i and δH gives

$$U_i = (A - RB^{-1})_{ij}^{-1} \left(\sum_q \frac{S_q \hat{r}_q^j}{\sigma_q^2} - B^{-1} \sum_q \sum_{q'} \frac{S_q r_q r_{q'}^j}{\sigma_q^2 \sigma_{q'}^2} \right), \quad (2)$$

and

$$\delta H = B^{-1} \sum_q \frac{S_q r_q - U_i r_q^i}{\sigma_q^2}, \quad (3)$$

where

$$A_{ij} = \sum_q \frac{\hat{r}_q^i \hat{r}_q^j}{\sigma_q^2}, \quad R_{ij} = \sum_q \sum_{q'} \frac{r_q^i r_{q'}^j}{\sigma_q^2 \sigma_{q'}^2}, \quad B = \sum_q \frac{r_q^2}{\sigma_q^2}. \quad (4)$$

S_q is related to the real peculiar velocity field $\mathbf{v}(\mathbf{r})$ by

$$S_q = v_i(\mathbf{r}_q) \hat{r}_q^i + \epsilon_q, \quad (5)$$

where ϵ_q is a gaussian random variable with an uncertainty σ_q . Therefore,

$$\delta H = \delta H^{(v)} + \delta H^{(\epsilon)}. \quad (6)$$

The first term is the noise-free contribution from density fluctuations

$$\delta H^{(v)} = \int d^3r W^i(\mathbf{r}) v_i(\mathbf{r}), \quad (7)$$

while the second term represents the contribution from noises in data.

The window function in eq. (7) is

$$W^i(\mathbf{r}) = \hat{r}^i B^{-1} \left\{ \sum_q \frac{r_q^i}{\sigma_q^2} \delta(\mathbf{r} - \mathbf{r}_q) - (A - RB^{-1})_{jl}^{-1} \left[\sum_q \frac{\hat{r}_q^j}{\sigma_q^2} \delta(\mathbf{r} - \mathbf{r}_q) - B^{-1} \sum_q \sum_{q'} \frac{r_q^i r_{q'}^j}{\sigma_q^2 \sigma_{q'}^2} \delta(\mathbf{r} - \mathbf{r}_q) \right] \sum_{q''} \frac{r_{q''}^l}{\sigma_{q''}^2} \right\}. \quad (8)$$

Its Fourier transformation is

$$W^i(\mathbf{k}) = \frac{B^{-1}}{(2\pi)^{3/2}} \left[\sum_q \frac{r_q^i}{\sigma_q^2} e^{i\mathbf{k} \cdot \mathbf{r}_q} - (A - RB^{-1})_{jl}^{-1} \left(\sum_q \frac{\hat{r}_q^i \hat{r}_q^j}{\sigma_q^2} e^{i\mathbf{k} \cdot \mathbf{r}_q} - B^{-1} \sum_q \sum_{q'} \frac{r_q^i r_{q'}^j}{\sigma_q^2 \sigma_{q'}^2} e^{i\mathbf{k} \cdot \mathbf{r}_q} \right) \sum_{q''} \frac{r_{q''}^l}{\sigma_{q''}^2} \right]. \quad (9)$$

The expected r.m.s. peculiar Hubble flow for the sample is

$$\left\langle \left(\frac{\delta H}{H_0} \right)^2 \right\rangle = \left\langle \left(\frac{\delta H^{(v)}}{H_0} \right)^2 \right\rangle + \left\langle \left(\frac{\delta H^{(\epsilon)}}{H_0} \right)^2 \right\rangle. \quad (10)$$

In linear theory,

$$\left\langle \left(\frac{\delta H^{(v)}}{H_0} \right)^2 \right\rangle = \Omega_0^{1.2} \int d^3k |W^i(\mathbf{k}) \hat{k}_i|^2 \frac{P(k)}{k^2}. \quad (11)$$

From eq. (5) and (6)

$$\begin{aligned} \left\langle \left(\frac{\delta H^{(\epsilon)}}{H_0} \right)^2 \right\rangle &= B^{-2} \left\langle \left[\sum_q \frac{\epsilon_q r_q}{\sigma_q^2} - (A - RB^{-1})_{ij}^{-1} \left(\sum_q \frac{\epsilon_q \hat{r}_q^j}{\sigma_q^2} - B^{-1} \sum_q \sum_{q'} \frac{\epsilon_q r_q r_{q'}^j}{\sigma_q^2 \sigma_{q'}^2} \right) \sum_{q''} \frac{r_{q''}^i}{\sigma_{q''}^2} \right] \right. \\ &\quad \left. \left[\sum_q \frac{\epsilon_q r_q}{\sigma_q^2} - (A - RB^{-1})_{lm}^{-1} \left(\sum_q \frac{\epsilon_q \hat{r}_q^m}{\sigma_q^2} - B^{-1} \sum_q \sum_{q'} \frac{\epsilon_q r_q r_{q'}^m}{\sigma_q^2 \sigma_{q'}^2} \right) \sum_{q''} \frac{r_{q''}^l}{\sigma_{q''}^2} \right] \right\rangle \\ &= B^{-1} + B^{-2} (A - RB^{-1})_{il}^{-1} R_{il}. \end{aligned} \quad (12)$$

Similarly,

$$\left\langle U_i^{(\epsilon)} U_j^{(\epsilon)} \right\rangle = (A - RB^{-1})_{ij}^{-1}. \quad (13)$$

If the true Hubble constant is uncertain and one wishes to investigate the variation of the Hubble flow within the sample, it is more appropriate to calculate the r.m.s. deviation of the Hubble flow of a subsample (with a rate H_2) from the Hubble flow of the whole sample (with a rate H_1). Thus

$$\begin{aligned} \left\langle \left(\frac{H_1 - H_2}{H_1} \right)^2 \right\rangle &= \left\langle \left(\frac{\delta H_1 + H_0 - \delta H_2 - H_0}{\delta H_1 + H_0} \right)^2 \right\rangle \approx \left\langle \left(\frac{\delta H_1}{H_0} \right)^2 + \left(\frac{\delta H_2}{H_0} \right)^2 - 2 \left(\frac{\delta H_1}{H_0} \right) \left(\frac{\delta H_2}{H_0} \right) \right\rangle \\ &= \left\langle \left(\frac{\delta H_1^{(v)}}{H_0} \right)^2 + \left(\frac{\delta H_2^{(v)}}{H_0} \right)^2 - 2 \left(\frac{\delta H_1^{(v)}}{H_0} \right) \left(\frac{\delta H_2^{(v)}}{H_0} \right) \right\rangle + \left\langle \left(\frac{\delta H_1^{(\epsilon)}}{H_0} \right)^2 + \left(\frac{\delta H_2^{(\epsilon)}}{H_0} \right)^2 - 2 \left(\frac{\delta H_1^{(\epsilon)}}{H_0} \right) \left(\frac{\delta H_2^{(\epsilon)}}{H_0} \right) \right\rangle, \end{aligned} \quad (14)$$

in which

$$\left\langle \left(\frac{\delta H_1^{(v)}}{H_0} \right) \left(\frac{\delta H_2^{(v)}}{H_0} \right) \right\rangle = \Omega_0^{1.2} \int d^3k |W_1^i(\mathbf{k}) \hat{k}_i W_2^{j*}(\mathbf{k}) \hat{k}_j| \frac{P(k)}{k^2} \quad (15)$$

and

$$\left\langle \left(\frac{\delta H_1^{(\epsilon)}}{H_0} \right) \left(\frac{\delta H_2^{(\epsilon)}}{H_0} \right) \right\rangle = \left\langle \left(\frac{\delta H_1^{(\epsilon)}}{H_0} \right)^2 \right\rangle. \quad (16)$$

Subscript “1” refers to quantities of the entire sample, and subscript “2” refers to those of the subsample.

3. Peculiar Hubble Flows in the Mark III catalogue

The Mark III catalogue (Willick et al. 1996) compiled the distances and redshifts to about 3000 galaxies, based on a template Tully-Fisher relation from the 36 clusters of Han & Mould (HM sample thereafter; Mould et al. 1991; Han and Mould 1992; Mould et al. 1993). Therefore the underlying Hubble expansion of the Mark III catalogue is defined by the HM sample. Table 1 lists the expected 1σ deviation of this underlying Hubble expansion rate H_1 from a global Hubble constant H_0 in various models, which is typically 2% to 3%. (In this and all the following calculations, clusters are assumed to have a gaussian random motion with a one-dimension dispersion of 300 km/sec, and all bulk motions are in reference to the Cosmic Microwave Background.) Figure 1(a) shows the variation of

Hubble flows in the HM sample with 2σ errorbars calculated from eqs. (2) to (16) vs. the depth R of subsamples. Each subsample is defined to be the N clusters closest to us, and R is the distance to the furthest cluster in the subsample. The errors of the variation are estimated from the second term of eq. (14). In particular, they include the contribution from uncertainties in determining bulk motions.

The figure shows significant positive variation within 40 to $60h^{-1}$ Mpc, translating into significant peculiar Hubble flows $H_2 - H_0$ given the smallness of $H_1 - H_0$. In figure 1(b) the variation is compared with 2σ expectations (with noise included) from theoretical models. Clearly the positive $(H_2 - H_1)/H_1$ found exceed the 2σ expectation from Λ CDM within $\sim 50h^{-1}$ Mpc, and exceed the 2σ expectations from CHDM, tCDM and Λ CDM within a wider range of subsamples. The variation at scales beyond $\sim 60h^{-1}$ Mpc, however, is perfectly consistent with all models at 2σ . Not only the variation of Hubble flows in the HM sample is extremely large, its bulk flow is also extremely large, at $(740 \pm 150 \text{ km/sec}, -267 \pm 134 \text{ km/sec}, -360 \pm 141 \text{ km/sec})$, compared to the noise-free 1σ Λ CDM expectation of 450 km/sec, or the noise-free 1σ tCDM expectation of 240 km/sec.

Neither the large variation of Hubble flows within $\sim 50h^{-1}$ Mpc nor the large bulk flows may be surprising. Figure 2(a) projects the positions of the inner 20 clusters (with a depth of $56h^{-1}$ Mpc) and the direction of their bulk motion on the sky. The four clusters at the lower right quadrant of the map, Telescopium at $24h^{-1}$ Mpc, OC3627 at $29h^{-1}$ Mpc, Pavo II at $37h^{-1}$ Mpc and OC3742 at $40h^{-1}$ Mpc, all show positive radial peculiar velocities. Since they lie roughly between us and the Great Attractor (GA, at ~ 4200 km/sec and in the direction of Hydro-Centaurus) (Lynden-Bell et al. 1988), their positive radial peculiar velocities may be the result of the gravitational pull of the GA. The clusters not in the general direction of the GA are further away from it and thus experience a smaller infall into the GA. As a result the subsample shows a large dipole in the direction

of the four clusters, plus a large positive monopolar flow. Even the entire HM sample with 36 clusters going up to 11,000 km/sec may show the influence of the GA, since from figure 2(b), a projection of the outer 16 clusters, no new cluster beyond the distance of the GA is added in the same direction of the four clusters, and almost all the outer 16 clusters are in the opposite hemisphere. To further show the effect of the four clusters, figure 3 plots $(H_2 - H_1)/H_1$ vs. the depth R of subsamples without the four clusters. The positive variation of Hubble flows at 40 to $60h^{-1}$ Mpc becomes much less significant and no longer exceeds the 2σ expectations of the four chosen models. The bulk motion of the remaining 32 clusters, however, is still large, at $(718 \pm 205 \text{ km/sec}, -216 \pm 140 \text{ km/sec}, -344 \pm 191 \text{ km/sec})$. Clearly, the HM sample is so sparse and inhomogeneous that spurious bulk motions and peculiar Hubble flows can result from structures such as the GA. The window functions $W^i(\mathbf{k})$ of the sample, on the other hand, may not fully reflect the density waves that give rise to these structures. Therefore the theoretical expectations based on these window functions can be too small. This also implies that the assumption that the HM sample well represents the true Hubble flow may rather be poor.

Similar calculations are done to the 277 galaxy groups of Mathewson et al. (1992) plus the 11 clusters of Willick (1991) (M+W sample thereafter) in the Mark III catalogue. The Aaronson et al. (1982) sample is too shallow to be included. The Mathewson et al. (1992) sample is much denser than the HM sample. It also goes up to $\sim 10,000$ km/sec, but only covers the southern sky. The Willick (1991) sample is too sparse, but it is not overly influenced by the GA (thanks to the large distances of its members from the GA). Figure 4(a) and 4(b) show the projections of the two samples on the sky. Figure 5(a) shows the variation $(H_2 - H_1)/H_1$ vs. the depth R of the subsamples of the M+W sample. Figure 5(b) shows the comparison between this variation and 2σ theoretical expectations (with noises included). The errors and noise-included theoretical expectations are calculated only for two representative points, because both quantities are monotonic, slowly varying

functions of R . It should also be pointed out that H_1 is the Hubble expansion rate of the HM sample because it provides the template Tully-Fisher relation for the Mark III catalogue. Therefore, instead of eq. (14),

$$\left\langle \left(\frac{H_2 - H_1}{H_1} \right)^2 \right\rangle \approx \left\langle \left(\frac{\delta H_1^{(v)}}{H_0} \right)^2 + \left(\frac{\delta H_2^{(v)}}{H_0} \right)^2 \right\rangle + \left\langle \left(\frac{\delta H_2^{(\epsilon)}}{H_0} \right)^2 \right\rangle. \quad (17)$$

The bulk motion of the M+W sample is relatively small, $(256 \pm 57 \text{ km/sec}, -343 \pm 61 \text{ km/sec}, 170 \pm 47 \text{ km/sec})$, while the noise-free expectation of Λ CDM is 670 km/sec, and that of t CDM is 361 km/sec.

There are two conclusions that can be drawn from figure 5: (1) the M+W sample shows Hubble flows that are 6% to 8% faster than the Hubble flow of the HM sample, with a $\sim 4\sigma$ significance; (2) the differences between the HM sample of 36 clusters and the combined M+W sample are consistent with theoretical expectations within 2σ , although the consistency is marginal for models that predict smaller peculiar velocities, such as t CDM and Λ CDM. But because of the poor sampling of the HM sample, whether the first conclusion implies a significant positive peculiar Hubble flow in the M+W sample is murky.

4. Peculiar Hubble Flows in the LP Catalogue

Lauer and Postman’s Abell cluster catalogue is a volume-limited sample that includes 119 Abell clusters with a redshift of $\lesssim 0.05$ and a galactic latitude above 13 degrees on both hemispheres. Its advantages over the HM sample are that it is volume-limited, more homogeneous, and deeper. As table 1 shows, its Hubble flow is expected to deviate from the global Hubble flow by only $\sim 1\%$. The standard candle of the catalogue is taken to be the luminosity L of the brightest cluster galaxies as a function of the second parameter α , the power index of L as a function of the aperture. The distance to a BCG is taken to be

its cosmological redshift ¹, i.e.,

$$r_q = (cz_q - U_i \hat{r}_q^i - \delta H_1 r_q) / H_0 = (cz_q - U_i \hat{r}_q^i) / (H_0 + \delta H_1) = (cz_q - U_i \hat{r}_q^i) / H_1, \quad (18)$$

where H_1 is the Hubble expansion rate defined by the entire sample. The estimated line-of-sight peculiar velocity is

$$S_q = cz 10^{0.4[M_*(\alpha_q) - M_q]/(2 - \alpha_q)}, \quad (19)$$

where M_q is the absolute magnitude of BCG q , and $M_*(\alpha_q)$ is the magnitude of the standard candle at $\alpha = \alpha_q$.

When applying eqs. (1) to (5) to the Lauer and Postman's catalogue, because M_q and α_q depend on r_q and thus U_i and δH , one has to iterate eqs. (1) to (5) to obtain a self-consistent result. Assuming no peculiar Hubble flow, Lauer and Postman (1994) got a bulk flow of the sample $U_x = 477 \pm 250$ km/sec, $U_y = -142 \pm 273$ km/sec, $U_z = 635 \pm 198$ km/sec relative to the rest frame of Cosmic Microwave Background Radiation. I first repeat Lauer and Postman's calculation, log-linearly interpolating table 3 of LP to find out the dependence of M_q and α_q on r_q . I get $U_x = 475 \pm 285$ km/sec, $U_y = -123 \pm 308$ km/sec, $U_z = 648 \pm 225$ km/sec, in very good agreement with LP's result. The errors are estimated from eq. (13), and verified by Monte Carlo simulations similar to those of LP.

I then consider the possibility of a non-zero $\delta H_1 = H_1 - H_0$. Since the distance scale is established within the sample itself, δH_1 of the sample is really not a variable but a well-defined value. In particular, one can always redefine the true Hubble expansion rate as $H'_0 = H_0 + \delta H_1 = H_1$ so that variation of Hubble flows in the sample can be investigated without referring to H_0 , just like in the previous section.

¹The cosmological distortion is small and omitted. See the discussion at the end of the section.

However, if one simply follows LP’s way of calibrating $M_*(\alpha)$, redefining δH_1 to zero cannot be done. This is due to the fact that the non-linear relation between S_q and $M_*(\alpha_q) - M_q$ always skews the distribution of S_q to the negative direction when $M_*(\alpha) - M_q$ is gaussian as calibrated by regression from the M - α distribution. An unphysical and negative $\delta H_1/H_0$ will always result for such a calibration, regardless of the value of H_0 . For the LP sample, $\delta H_1/H_0 \approx -1.2\%$.

One way to get rid of the unphysical δH_1 and set δH_1 to 0 is to calibrate $M_*(\alpha)$ by regression from the (S_q/cz_q) - α distribution (my calculation shows that the resulting S_q/cz_q residual is gaussian with a standard deviation of 0.166 at 6% C.L.). The resultant bulk motion of the sample is then $U_x = 528 \pm 285$ km/sec, $U_y = -272 \pm 311$ km/sec, and $U_z = 607 \pm 225$ km/sec. U_y being the most uncertain component, shows the biggest difference from the LP result. The resultant δH_1 converges to zero to a high precision.

Once $M_*(\alpha)$ is calibrated, one can calculate the variation of Hubble flows within the sample. Figure 6 shows the variation with 2σ error bars vs. the depth of the LP subsamples, and its comparison with noise-included 2σ theoretical expectations. The error bars are once again calculated using eq. (14). They are checked with Monte Carlo simulations and consistency is found. The figure shows evidence for negative $(H_2 - H_1)/H_1$ within a radius of $\sim 60h^{-1}$ Mpc at the 2σ level, indicating a negative peculiar Hubble flow at this scale given the tiny deviation of H_1 from H_0 (table 1). But before the negative peculiar Hubble flow is explained with overdensities, possible systematic effects have to be considered first, as in the case of the Mark III catalogue.

Possible biases in the analysis have been discussed extensively in LP (1994). Among them are several radially dependent biases that affect the calculation of peculiar Hubble flows. The first is the selection bias. A luminosity-limited catalogue introduces an artificial Hubble outflow due to the missing low brightness galaxies. This is apparently not the case

with LP’s catalogue because it is volume-limited. Secondly, the estimated peculiar velocity depends on the deceleration parameter q_0 assumed. But for a catalogue extending only to $z \approx 0.05$, the estimated radial peculiar velocity is only changed by $\lesssim 1\%$ if the true q_0 is changed by 0.5. Since the standard candle in LP’s catalogue is established internally, influenced mostly by outlying clusters, the effect on the peculiar Hubble flow of inner clusters is $\lesssim 1\%$. A third bias comes from the random peculiar velocities of BCGs due to local non-linearities which tend to scatter more BCGs to lower measured redshifts than to higher redshifts. But given a typical value of this random radial velocity of 300 km/sec, the velocity bias introduced on the 6000 km/sec scale is only 0.5% (LP). Another important concern is whether the BCGs of the inner Abell clusters belong statistically to the same population of the entire BCG sample. Table 2 lists the statistical properties of the inner BCG subsamples and the entire BCG sample. They are consistent statistically.

But a problem arises when one tests whether the detection is dominated by a small number of clusters. Figure 7 projects the inner 11 clusters (which show the negative peculiar Hubble flow at 2σ) on the sky. The subsample is very inhomogeneous, due to its small size. In particular, there is one (and only one) cluster (A262) that roughly aligns with the bulk motion of the subsample. Since the bulk motion and the peculiar Hubble flow are not determined independently, the cluster will certainly play a dominate role in determining both quantities. If the clusters is excluded, as shown in the small window of figure 6(a), no variation of Hubble flows is found at the 2σ level. Therefore, once again, the poor sampling of the subsample renders its result susceptible to systematic effects such as small scale structures, unrepresented structures, small number statistics, etc.

5. Peculiar Hubble Flows in a Type Ia supernova sample

Type Ia supernova samples probe much deeper ($\sim 10^3$ Mpc) and have very precise distance measurements ($\sim 5\%$). Therefore, even for the 20 Type Ia SNe compiled by Riess et al. (1996), table 1 shows that the expected 1σ deviation of its Hubble flow from the global one is only $\sim 1\%$. But because of its extremely sparse sampling, its usefulness to investigations of peculiar Hubble flows is significantly compromised. Figure 8 shows the variation of Hubble flows in the Type Ia SN sample. No significant variation is found at any scale because of its sparse sampling. Since the errors in the plot roughly scale as the inverse square root of the number of objects, as more Type Ia supernovae are being observed with good time coverage, they can certainly provide more accurate results at a wide range of depth in the future.

6. Discussion

After investigating the variation of Hubble flows in the Mark III catalogue (Willick et al. 1996), the Lauer and Postman (1994) sample, and the Type Ia supernova sample of Riess et al. (1996), there are three conclusions that I would like to draw:

- Some significant peculiar Hubble flows are found at face value in the Mark III catalogue and at 50 to $60h^{-1}$ Mpc in the Lauer and Postman sample. They at face value disfavor cosmological models predicting smaller peculiar velocities, such as Λ CDM. However, because of the sparse and inhomogeneous sampling of the HM cluster sample on which the Hubble expansion of the Mark III catalogue is defined, and of the sparse sampling of the Lauer and Postman sample within the $60h^{-1}$ Mpc scale, the peculiar Hubble flows found are dominated by statistics of a small number of clusters, and thus cannot be trusted at their face value. Further supporting this is

the fact that the peculiar Hubble flows found in the different samples are inconsistent: the HM sample shows positive peculiar Hubble flows below $\sim 70h^{-1}$ Mpc, while the M+W sample shows positive peculiar Hubble flows up to $\sim 100h^{-1}$ Mpc; the LP sample, on the other hand, shows negative peculiar Hubble flows at $\sim 60h^{-1}$ Mpc.

- All samples agree on that the Hubble flows beyond a depth of $\sim 80h^{-1}$ Mpc are true to the global Hubble flow to better than 10% at $\sim 2\sigma$ C.L. In each sample, the variation of Hubble flows at a depth of $\geq 80h^{-1}$ Mpc, with respect to the Hubble flow of entire sample, are found to be $< 10\%$ at 2σ . This combined with the small expected deviation (1 to 3% at 1σ , see table 1) of the Hubble flow defined by each sample from the global Hubble flow, ensures that the Hubble flows at $\geq 80h^{-1}$ Mpc conform to the global Hubble flow to better than 10% at $\sim 2\sigma$ C.L. At depth below $80h^{-1}$ Mpc, the limits on peculiar Hubble flows are somewhat weaker. For example, at the depth of the Coma cluster ($70h^{-1}$ Mpc), while the Lauer and Postman sample indicates a less than 15% deviation from the true Hubble flow at 2σ , all others imply a stronger limit—less than 10% deviation from the true Hubble flow at 2σ C.L.
- Since the peculiar Hubble flows detected in the samples are not trustworthy at the moment, so is its comparison with theoretical predictions. However, this is not to say that any such comparison will be useless, because the major limiting factor to a meaningful comparison between models and observations is neither our uncertain knowledge of the Hubble constant, nor the uncertainties in the distance scales, but rather the imperfect sampling of our local universe. The errors achieved by the current Mark III catalogue are comparable ($\sim 1\%$ at $\sim 60h^{-1}$ Mpc) to the differences in noise-free model predictions. Therefore, If samples were improved to yield a consistent and reliable detections of peculiar Hubble flows, measurements of peculiar Hubble flows would certainly be able to test models.

While the first and second conclusions represent an attempt to map out Hubble flows in our local universe, which are extremely interesting in their own rights, the third conclusion aims at finding a potential tool to gain further information on cosmological parameters, and therefore warrants a further discussion. Since errors in the question are roughly proportional to the square root of the number of objects in a sample, if a deep ($> 100h^{-1}$ Mpc), homogeneous (covering all directions and depth) and dense (at least 10^2 clusters to yield a template Tully-Fisher relation and $\sim 10^3$ clusters and groups within $100h^{-1}$ Mpc to be sampled) sample can be assembled with the Tully-Fisher relation, a reliable detection of peculiar Hubble flows may be made with errors at the $\lesssim 1\%$ level, sufficient to test many models at the $\sim 60h^{-1}$ Mpc scale. Or, if more than 100 Type Ia supernovae are adequately observed as distance indicators, the errors of their measured peculiar Hubble flows can be cut by more than 1/2 from the current level to $\sim 1\%$. Better yet, if the Tully-Fisher relation is calibrated with Type Ia supernovae, by observations of Type Ia supernovae in clusters, not only the overall Hubble flow of the Tully-Fisher sample will represent the global Hubble flow better, the two samples can be combined to yield better statistics.

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Table 1. Parameters of four chosen models*and their predicted 1σ deviations from a global Hubble flow for the Han & Mould sample (HM), the Lauer & Postman sample (LP), and the Type Ia supernova sample (SN).

	sCDM	tCDM	Λ CDM	CHDM
Ω_{tot}	1	1	1	1
Ω_m	1	1	0.3	1
Ω_b	0.06	0.06	0.03	0.06
Ω_ν	0	0	0	0.2
h	0.5	0.5	0.7	0.5
n	1	0.7	1	0.9
$\left\langle \left(\frac{H_1 - H_0}{H_0} \right)^2 \right\rangle^{1/2}_{\text{HM}}$	3.1%	1.8%	2.1%	2.4%
$\left\langle \left(\frac{H_1 - H_0}{H_0} \right)^2 \right\rangle^{1/2}_{\text{LP}}$	1.2%	0.7%	0.9%	0.9%
$\left\langle \left(\frac{H_1 - H_0}{H_0} \right)^2 \right\rangle^{1/2}_{\text{SN}}$	2.0%	1.0%	1.1%	1.3%

*The power spectra of the models are from Bardeen et al. (1986), Bunn and White (1996), Ma (1996) and Sugiyama (1995).

Table 2. Comparison of statistical properties in the LP sample.

N ^a	$\langle\alpha\rangle$ ^b	$\langle M_q - M_*(\alpha_q)\rangle$ ^c	$\sigma_{M_q - M_*(\alpha_q)}$ ^d	P_{KS} ^e
8	0.61	0.141	0.255	83%
9	0.61	0.113	0.253	71%
10	0.59	0.109	0.239	84%
11	0.59	0.106	0.227	93%
119	0.57	0	0.245	39%

^aThe number of clusters in the (sub)sample.

^bThe average α in the (sub)sample.

^cThe average $M_q - M_*(\alpha_q)$ in the (sub)sample.

^dThe standard deviation of $M_q - M_*(\alpha_q)$.

^eThe confidence level that $M_q - M_*(\alpha_q)$ is gaussian according to a K-S test.

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Fig. 1.— (a) The variation of Hubble flows in the HM sample with 2σ error bars, as a function of the depth of subsamples. (b) The crosses show the variation of Hubble flows in the HM sample. Curves are the 2σ expectations (with noises included) from models. Model expectations for subsamples are connected for a simpler presentation.

Fig. 2.— (a) Projection of the inner 20 clusters of the HM sample and their bulk motion in galactic coordinates. Solid circles denote clusters with negative radial peculiar velocities, open circles denote clusters with positive radial peculiar velocities. The cross is the direction of the bulk motion. (b) Projection of the outer 16 clusters of the HM sample in galactic coordinates. The dotted line is the equator of the sky when the bulk motion of the inner 20 clusters is chosen as a pole.

Fig. 3.— The variation of Hubble flows in the HM sample less four clusters in the lower right quadrant of figure 2(a), with 2σ error bars.

Fig. 4.— (a) Projection of 277 groups of the Mathewson et al. (1992) sample in galactic coordinates. (b) Projection of 11 clusters of the Willick (1991) sample in galactic coordinates. Legends are the same as in figure 2.

Fig. 5.— Similar to figure 1 for the M+W sample. Errors and theoretical expectations are calculated only for one subsample and the entire sample.

Fig. 6.— Similar to figure 1 for the Lauer and Postman (1994) sample. Plot inside the window of (a) shows the variation of Hubble flows when cluster A262 is excluded.

Fig. 7.— Projection of the inner 11 Abell clusters in the Lauer and Postman sample and their bulk motion in galactic coordinates. Legends are the same as in figure 2.

Fig. 8.— Similar to figure 1(a) for the sample of 20 Type Ia supernovae.















